

MIDTERM REVIEW

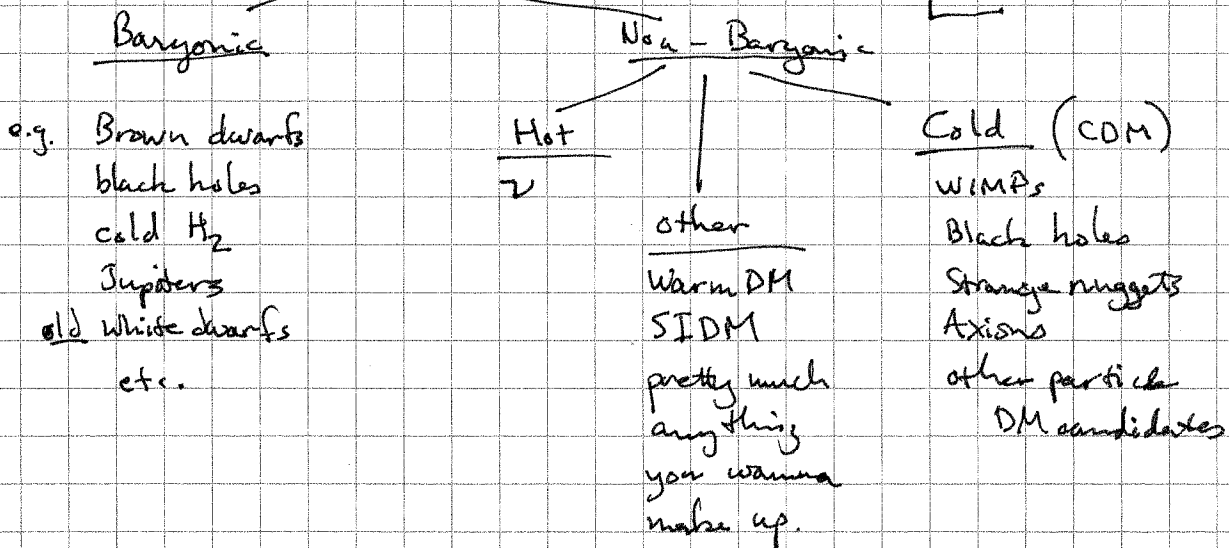
Observational Evidence for mass discrepancies

- Oort discrepancy in solar neighborhood
- flat rotation curves of spiral galaxies
- clusters of galaxies
- velocity dispersions
- hydrostatic equilibrium of X-ray gas
- gravitational lensing of background galaxies
- large scale structure
- $\Omega_m > \Omega_b$

Early indications included

- Oort (1932) - factor of ~ 2 discrepancy in solar neighborhood
- Zwicky (1933) - factor of ~ 100 discrepancy in clusters
- Ostriker & Peebles (1973) - factor of ~ 10 discrepancy in bar instability in disks

Dark Matter Candidates



WIMPs favored because

- $\Omega_m > \Omega_b$
- LSS formation

Dark matter is the usual inference; the discrepancy might also indicate a change in dynamical laws

Virial Theorem

Can be derived from stationary moment of inertial tensor

$$\text{boils down to } 2\langle K \rangle + \langle W \rangle = 0$$

Kinetic E Potential Energy

for N particles of equal mass m such that $M = Nm$,

$$M = \frac{2\sigma^2 R_{\text{rms}}}{G} \quad \text{where the harmonic radius } R_{\text{rms}}$$

is usually approximated as $R_{\text{rms}} \approx 1.25 R_d$

Vertical Force (Oort : restoring force to disk)

$$K_z = - \frac{\partial \Phi}{\partial z} = \frac{1}{2} \frac{\partial (v_\sigma^2)}{\partial z}$$

where $v(z)$ is the vertical profile of tracer population

locally this boils down to

$$\sigma_z^2 = 2\pi G \Sigma z_0$$

e.g.

$$v(z) = v_0 e^{-z/z_0}$$

Disk Stability

LOCAL: Toomre Q :

$$Q = \frac{\sigma R K}{3.36 G \Sigma} \quad \text{locally stable if } Q \gg 1$$

GLOBAL:

$$X_m = \frac{K^2 R}{2\pi m G \Sigma} \quad \text{higher surface densities less stable}$$

Ostriker & Peebles:

$$t \lesssim 0.14 \quad \text{where } t = \frac{T}{|W|}$$

with

$$K = T + \frac{1}{2}\Pi$$

T = rotational kinetic energy
 Π = kinetic energy in random motions

Exponential Disks

2D face-on surface brightness profile

$$\Sigma(R) = \Sigma_0 e^{-R/R_d}$$

Σ_0 = central surface density
 R_d = scale length

This integrates to a total luminosity $L = 2\pi \Sigma_0 R_d^2$

The enclosed luminosity is a simple fun of ~~the~~ scale length

$$L(<x) = 2\pi \Sigma_0 R_d^2 \left[1 - (1+x)e^{-x} \right] \quad \text{where } x = \frac{R}{R_d}$$

in 3D one can have a "double exponential" model

$$\rho = \rho_0 e^{-R/R_d} e^{-z/z_0}$$

For general light profiles we can fit the

Sersic profile

$$\Sigma(R) = \Sigma_e e^{-b_n \left[\left(\frac{R}{R_e} \right)^n - 1 \right]}$$

which reduces to the exponential form for $n=1$
and is equivalent to the de Vaucouleurs profile for $n=4$

Potential - Density Pairs

Poisson Eqn $\nabla^2 \Phi = 4\pi G \rho$

has an analytic solution for a handful of Φ - ρ pairs.

\therefore It helps to know ∇^2 in the right coordinate system

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Phase space $f(\vec{x}, \vec{v}, t)$: the distribution function

describes the density distribution of particles in both configuration space (x, y, z) and momentum (v_x, v_y, v_z) .

For a stationary (stable) system of widely separated stars, we have the collisionless Boltzmann eqn

$$\frac{\partial f}{\partial t} + \vec{v} \cdot \vec{\nabla} f - \vec{\nabla} \Phi \cdot \frac{\partial f}{\partial \vec{v}} = 0$$

Note that the density in configuration space ν is just the integral over f in momentum:

$$\nu = \int f d^3v$$

We usually use this to break the collisionless Boltzmann eqn into the

Jean's eqns

$$\frac{\partial \nu}{\partial t} + \frac{\partial (\nu \bar{v}_i)}{\partial x_i} = 0$$

$$\frac{\partial (\nu \bar{v}_j)}{\partial t} + \frac{\partial (\nu \bar{v}_i \bar{v}_j)}{\partial x_i} + \nu \frac{\partial \Phi}{\partial x_j} = 0$$

$$\nu \frac{\partial v_j}{\partial t} + \nu \bar{v}_i \frac{\partial \bar{v}_j}{\partial x_i} = -\nu \frac{\partial \Phi}{\partial x_j} - \frac{\partial (\nu \sigma_{ij}^2)}{\partial x_i}$$

where we have the "moments" of f

$$\nu = \int f d^3v$$

$$\bar{v}_i \bar{v}_j = \frac{1}{\nu} \int v_i v_j f d^3v$$

$$\bar{v}_i = \frac{1}{\nu} \int v_i f d^3v$$

$$\sigma_{ij}^2 = \bar{v}_i v_j - \bar{v}_i \bar{v}_j$$

Energy & Angular momentum

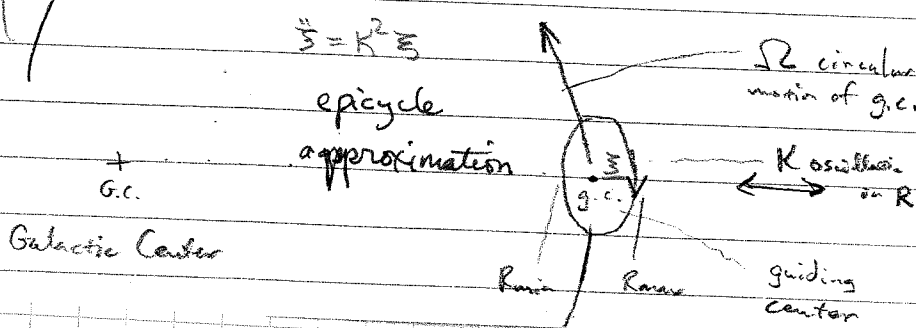
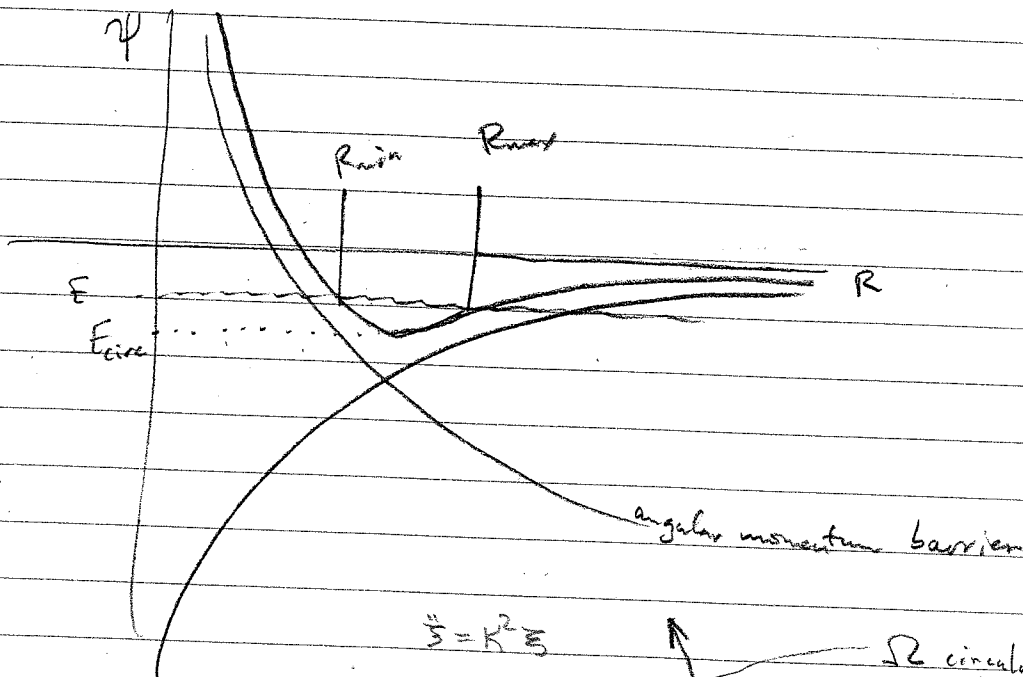
$$E = \frac{1}{2}(V_R^2 + V_\phi^2 + V_z^2) + \Phi(R, z) \quad \text{energy per unit mass}$$

in cylindrical coordinates

$$J_z = R V_\phi \quad \text{angular momentum per unit mass}$$

Effective Potential

$$\Psi(R, z) = \Phi(R, z) + \frac{J_z^2}{2R^2}$$



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Galactic constants

R_0 - distance to Galactic Center

V_0 - circular velocity of LSR (sometimes called Θ_0)

$$\Omega_0 = \frac{V_0}{R_0} \text{ - orbital frequency}$$

0 subscript denotes solar location orbital period
 $P = \frac{2\pi}{\Omega}$

Oort "constants"

$$A = -\frac{1}{2} \left[R \frac{d\Omega}{dR} \right]_{R_0} = \frac{1}{2} \left(\frac{V}{R} - \frac{dV}{dR} \right)_{R_0} \quad \text{SHEAR}$$

$$B = \frac{1}{2} \left(\frac{V}{R} + \frac{dV}{dR} \right)_{R_0} \quad \text{VORTICITY}$$

due to
angular
momentum
gradient

NOTE : $\Omega = A - B$

$$-\frac{dV}{dR} = A + B$$

epicyclic frequency : $K^2 = -4B\Omega$

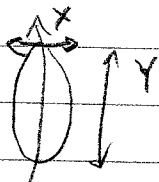
frequencies $\omega_z > K > \Omega$ so orbits not closed

more vertical oscillations than radial than complete orbits

$$\omega_z \approx 48$$

$$K \approx 37$$

$$\Omega \approx 30 \text{ km s}^{-1} \text{ kpc}^{-1} \text{ in solar neighborhood}$$



size of ellipsoid

$$\frac{Y}{X} = \frac{2\Omega}{K} = \frac{\sigma_x}{\sigma_y}$$

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timescales

crossing time

$$t_c = \frac{2R}{V}$$

dynamical time

$$t_d = \sqrt{\frac{3\pi}{16G\rho}}$$

relaxation time

$$\frac{t_r}{t_c} = \frac{N}{48f^2} \approx \frac{N}{6 \ln(N/2)}$$

3 LAWS of GALACTIC ROTATION

1. Flat rotation curves

The rotation curves of rotating galaxies tend to approach an approximately constant velocity that persists to indefinitely large radii.

2. Baryonic Tully-Fisher Relation: $M_b = AV_f^4$

The total baryonic mass of a rotating galaxy scales as the fourth power of its flat rotation velocity.

3. Mass Discrepancy - Surface Density relation

The amplitude of the mass discrepancy scales roughly as $\Sigma_b^{1/2}$.
(Holds both globally and locally)

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Halo models

pseudo-isothermal

empirically motivated

characterized by
core radius R_c
flat velocity V_{∞}

$$\rho_{iso}(r) = \frac{\rho_0}{1 + (r/R_c)^2}$$

$$V(r) = V_{\infty} \sqrt{1 - \frac{R_c}{R} \tan^{-1}\left(\frac{R}{R_c}\right)}$$

$$V_{\infty} = \sqrt{4\pi G \rho_0 R_c^2}$$

NFW

derived from simulations

characterized by

$$\rho_{NFW}(r) = \frac{4\rho_s}{(r/r_s)(1+r/r_s)^2}$$

c
 $R_{200}/V_{200}/M_{200}$

$$V(r) = V_{200} \sqrt{\frac{\ln(1+cx) - \frac{cx}{1+cx}}{x \left[\ln(1+c) - \frac{c}{1+c} \right]}}$$

$$x = \frac{r}{R_{200}}$$

~~$$x = \frac{r}{R_{200}}$$~~

$$c = \frac{R_{200}}{r_s}$$

$$V_{200} = h R_{200}$$

where $h = \frac{H_0}{100}$

in km/s in kpc

$$M_{200} = \frac{4\pi}{3} (200)^3 R_{200}^3$$

mention Einasto
Barkert

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