



Testing dark matter models and modified gravity theories with spatially resolved HI observations

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Understanding the nature of dark matter (DM) is one of the most pressing questions of modern physics. The flat rotation curves of galaxies, especially the extended H I rotation curves from radio observations, played a key historical role to establish the DM problem and continue to be a major tool to test different galaxy formation models, particle DM models, and modified gravity theories.

The square kilometre array (SKA) observatory will revolutionize the study of H I rotation curves thanks to its unprecedented angular resolution, sensitivity, and survey speed. In combination with optical and near-infrared ancillary data, H I surveys with the SKA-Mid will allow us to (1) obtain H I data at spatial resolutions of $1'' - 2''$, so study the inner galaxy dynamics with unprecedented details (e.g., the cusp-core problem); (2) move from representative samples of a few hundreds of objects to statistical samples of tens of thousands of objects, so study the dynamical scaling laws of galaxies and their DM properties with unprecedented statistical power across different cosmic environments; (3) move from galaxies at $z \simeq 0$ to galaxies at $z \simeq 1$, probing DM halos at earlier cosmic epochs and exploring the redshift evolution of the dynamical laws across ~ 8 Gyrs.

The massive SKA data, however, will pose important technical challenges for the derivation of H I rotation curves with automated kinematic software. Developing such kinematic software in the SKA era will be crucial to reach major scientific goals, such as discriminating between different solutions to the DM problem.

1 Introduction

The nature of dark matter (DM) is one of the biggest mysteries of the Universe. DM is needed to explain a variety of astrophysical phenomena, such as the internal dynamics of galaxies and galaxy systems (groups and clusters), the strength of gravitational lensing, the power spectrum of the cosmic microwave background (CMB), and the dynamical evolution of the Universe as a whole (for historical perspectives, see [Sanders, 2010](#); [Bertone and Hooper, 2018](#)). DM is thought to be a new elementary particle, beyond the standard model of particle physics, but has eluded clear detections so far (e.g., [Billard et al., 2022](#)). Nonbaryonic DM is a major ingredient of the Λ Cold Dark Matter (Λ CDM) cosmological model and is essential to promote structure formation, starting from the tiny density perturbations imprinted on the CMB up to present-day galaxies. Alternatively, the DM effect may indicate the need to revise Einstein’s theory of general relativity and its Newtonian limit. In this context, Modified Newtonian Dynamics (MOND, [Milgrom, 1983a](#)) and its relativistic extensions ([Skordis and Złóśnik, 2021](#); [Blanchet and Skordis, 2024](#)) appear to be most promising. Without a doubt, understanding the nature of DM is one of the major challenges of modern physics.

The flat rotation curves of disk galaxies played a key historical role in the establishment of the DM problem ([Rubin et al., 1978](#); [Bosma, 1978](#)). A key complementary role was also played by optical and near-infrared (NIR) surface photometry, providing the stellar gravitational contribution. In fact, the proper computation of the baryonic gravitational field of a disk-like system ([Kalnajs, 1983](#); [Casertano, 1983](#); [Kent, 1986](#)) showed that the $H\alpha$ rotation curves from optical spectroscopy were *not* extended enough to unambiguously probe the DM effect. The decisive evidence came from radio observations of the 21-cm line of atomic hydrogen (HI). The HI disks of galaxies are generally twice as extended as their stellar components, so they are uniquely suited to probe the outermost regions where the DM effect is dominant ([van Albada and Sancisi, 1986](#); [Kent, 1987](#)).

Over the past 40 years, rotation curves have been derived for hundreds of nearby galaxies (e.g., [Begeman et al., 1991](#); [Persic et al., 1996](#); [Verheijen and Sancisi, 2001](#); [Swaters et al., 2002](#); [de Blok et al., 2008](#)). Ideally, one combines $H\alpha$ and/or CO data at high angular resolutions ($\sim 1'' - 2''$) in the inner parts with HI data at lower angular resolutions ($\sim 10'' - 30''$) in the outer regions (e.g., [de Blok and Bosma, 2002](#); [Noordermeer et al., 2005](#); [Lelli et al., 2016b](#); [Shelest and Lelli, 2020](#)). In addition, NIR surface photometry provides the best possible tracer of the stellar mass distribution (e.g., [Verheijen and Sancisi, 2001](#); [McGaugh and Schombert, 2014](#); [Schombert et al., 2019](#)), so it can be used in concert with HI and CO surface density maps to construct full mass models, separating the gravitational contributions of stars, gas, and DM (see Fig. 1 for an example).

The Spitzer Photometry and Accurate Rotation Curves (SPARC) database has been a major advance in this effort, collecting HI+ $H\alpha$ rotation curves from the literature (see [Lelli et al., 2016b](#), references in) and complementing them with NIR surface photometry ([Schombert and McGaugh, 2014](#)). The SPARC sample consists of 175 galaxies at $z \simeq 0$ and has been used to test different Λ CDM models of galaxy formation ([Katz et al., 2017](#); [Li et al., 2020](#)) as well as alternative particle DM models, such as self-interacting DM ([Ren et al., 2019](#)), fuzzy DM ([Bar et al., 2022](#)), and superfluid DM ([Mistele et al., 2022](#)). The SPARC data have also been used to test modified gravity theories, such as MOND ([Li et al., 2018](#)), massive gravity ([Panpanich and Burikham, 2018](#)), $f(R)$ gravity ([Naik et al., 2019](#)), conformal gravity ([Li and Modesto, 2020](#)), and others. However, the SPARC sample

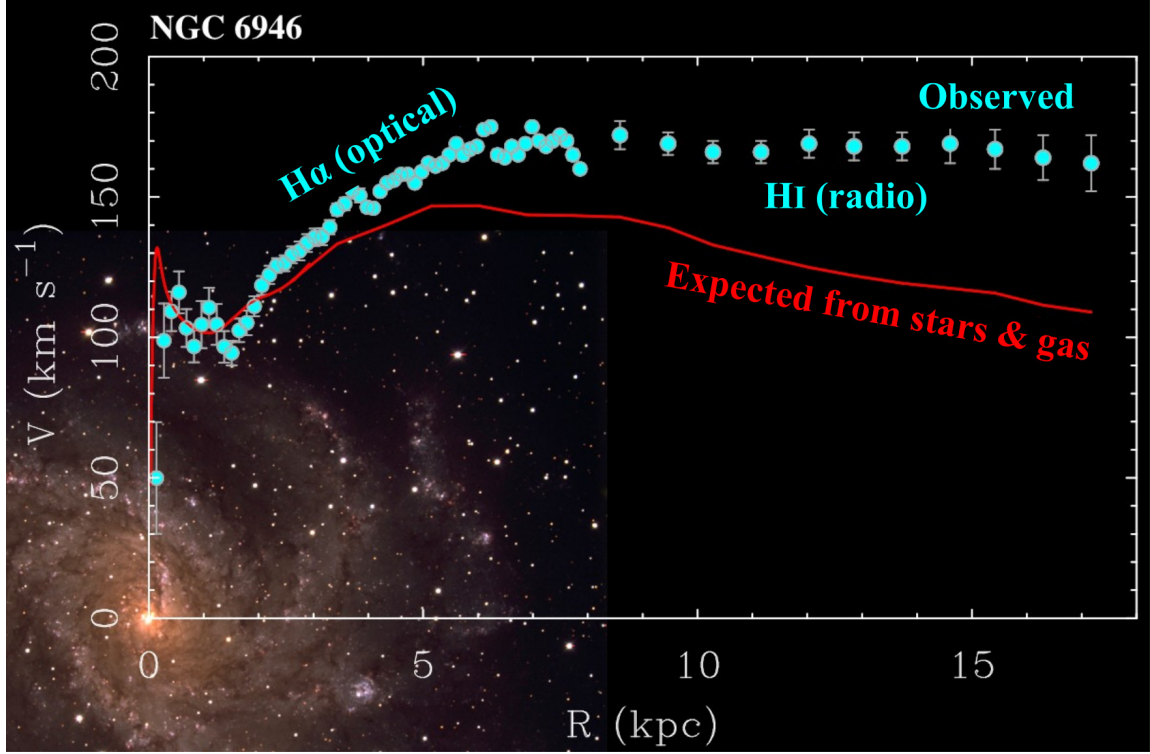


Figure 1: A classic evidence for the DM effect: the observed rotation curve of the nearby galaxy UGC6946 (cyan dots) differs from that expected from the distribution of baryons (red line). The red line is computed solving the Poisson’s equation for the observed distribution of stars (using the NIR image in the background) and gas (using the total HI map, not shown). The H α rotation curve ($R < 8$ kpc) from optical spectroscopy (Daigle et al., 2006) implies only a moderate mass discrepancy; the DM effect is most evident only at the larger radii probed by the HI rotation curve from radio interferometry (Boomsma et al., 2008). The figure has been created using data from the SPARC database (Lelli et al., 2016b).

remains too small and heterogeneous to firmly distinguish between these different theories. In addition, it lacks very low-mass galaxies (with baryonic masses smaller than $10^8 M_{\odot}$), which are probed by smaller complementary samples, such as LITTLE-THINGS (Oh et al., 2015; Iorio et al., 2017), SHIELD (McNichols et al., 2016), and the compilation of Mancera Piña et al. (2025).

To distinguish among different DM models and/or modified gravity theories, we need much larger samples (10-100 thousands of galaxies) with both HI and NIR data, probing galaxies in different cosmic environments. The first steps in this direction are represented by the WALLABY survey with the Australian SKA pathfinder (ASKAP, Westmeier et al., 2022; Deg et al., 2022; Murugesan et al., 2024), the AWES survey with APERTIF (Adams et al., 2022), and the upcoming BIG-SPARC database (Haubner et al., 2024). These projects will provide spatially resolved HI data for a few thousand galaxies, increasing current samples by about one order of magnitude. Another missing piece of evidence is probing the HI dynamics at $z \gtrsim 0$ in order to trace the cosmic evolution of the baryon-DM connection and/or the possible variation of the breakdown of modified gravity theories. The first steps in this direction are the MIGHTEE-HI survey (Jarvis et al., 2025; Vărașteanu et al., 2025) and the LADUMA survey (Smith, 2024) with the SKA-MID precursor MeerKAT.

The Square Kilometre Array (SKA) observatory, especially SKA-Mid in South Africa, will revolutionize the study of HI rotation curves due to its unprecedented angular resolution, sensitivity, and survey speed. This chapter describes the main prospects expected from the SKA-Mid Array Assembly 4 (AA4). We start with a description of the main open questions in galaxy dynamics, namely the origin of the dynamical scaling laws (Sect. 2.1) and some long-standing “small-scales” Λ CDM problems (Sect. 2.2). Next, we discuss the SKA prospects to obtain HI observations (1) at angular resolutions of $\sim 1'' - 2''$ comparable to those of $H\alpha$ and CO surveys (Sect. 3.1), (2) for statistical galaxy samples of tens of thousands of spatially resolved galaxies (Sect. 3.2), and (3) at redshifts $z \simeq 0.1 - 1.0$ that have been inaccessible in HI emission so far. We conclude with a discussion of next-generation tools to model HI kinematics in the SKA era (Sect. 4).

2 Open questions

2.1 The baryon-DM coupling and the dynamical laws of galaxies

A fundamental but somewhat underappreciated piece of evidence in the DM puzzle is that the dark mass and the luminous mass of galaxies appear tightly coupled to each other. Empirically, the observed dynamics, which is generally dominated by DM, are closely related to the distribution of baryons (e.g., [Swaters et al., 2012](#)). One of the pioneers of HI kinematics in galaxies, Renzo Sancisi, summarized the evidence with a concise statement, which is often referred to as Renzo’s Rule or Sancisi’s Law: “For any feature in the luminosity profile there is a corresponding feature in the rotation curve and vice versa” ([Sancisi, 2004](#)). Another highlighting quote from [Sancisi \(2004\)](#) is the following one: “The unavoidable conclusion from the observed correspondence between the shapes of the rotation curves and those of the luminosity profiles is that the gravitational potential is strongly correlated with the distribution of luminous matter”.

The empirical baryon-DM coupling can be quantified by a set of “dynamical laws” that connect a baryonic property (inferred from the observed distribution of stars and gas) and a dynamical one (inferred from the shape of the rotation curve), making no assumption about the nature of DM. Figure 2 shows three of the main dynamical scaling laws of galaxies (see [Lelli, 2022](#), for details): (1) the baryonic Tully-Fisher relation (BTFR) that links the total baryonic mass to the “flat” rotation velocity along the outer part of the rotation curve ([McGaugh, 2005](#); [Lelli et al., 2016a, 2019](#)); (2) the central density relation (CDR) that links the central baryonic surface density to the central dynamical-mass surface density inferred from the inner rising part of the rotation curve ([Lelli et al., 2013, 2016c](#); [Swaters et al., 2014](#)); and (3) the radial acceleration relation (RAR) that links the baryonic gravitational field measured from the baryonic distribution to the observed centripetal acceleration at each radius ([McGaugh et al., 2016](#); [Lelli et al., 2017](#)). These dynamical laws are extremely tight by the usual standards of extragalactic astronomy: the observed scatter around the mean relations is consistent with observational uncertainties, implying that any intrinsic scatter must be tiny ([Li et al., 2018](#); [Desmond, 2023](#)). The properties of these dynamical laws – in particular their intrinsic tightness – represent one of the major “small-scale challenges” for Λ CDM models of galaxy formation ([Bullock and Boylan-Kolchin, 2017](#)).

The dynamical laws of galaxies raised a lively debate in the community because they can be interpreted in drastically different ways. In a particle DM scenario (irrespective of CDM, self-

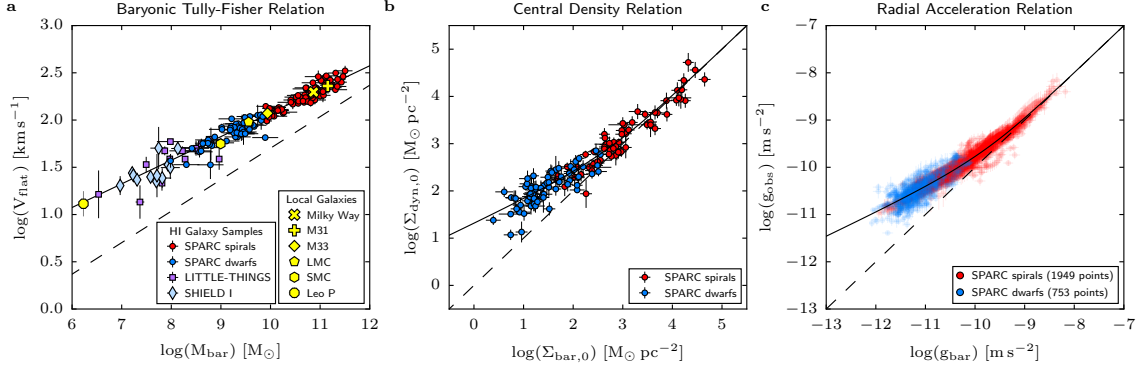


Figure 2: The dynamical scaling laws of nearby galaxies (adapted from Lelli, 2022). **Panel a:** The baryonic Tully-Fisher relation (BTFR). The dashed line shows the simplest Λ CDM prediction, with slope of 3 and intercept set by the cosmological ratio $\Omega_{\text{bar}}/\Omega_{\text{CDM}}$ (see, e.g., McGaugh, 2012). The solid line shows a fit fixing the slope to 4, as predicted by MOND. **Panel b:** The central density relation (CDR). The dashed line corresponds to the 1:1 line. The solid line shows a one-parameter fit using the MOND interpolation function. **Panel c:** the radial acceleration relation (RAR). The dashed line corresponds to the 1:1 line. The solid line shows a one-parameter fit using the MOND interpolation function. In the BTFR and CDR, each point represents a single galaxy. In the RAR, each point represents a spatially-resolved measurement at different radii in 153 different galaxies. In all cases, the observed scatter is comparable to that expected from observational uncertainties, leaving little room for intrinsic scatter.

interacting DM, fuzzy DM, and so on), the physical processes driving galaxy evolution inside DM halos must conspire to produce these scaling relations at $z = 0$, which would then represent an “attractor solution” for baryonic physics (e.g., Di Cintio and Lelli, 2016; Ludlow et al., 2017). In a modified gravity scenario, instead, these scaling relations may represent new “Laws of Nature” akin to Kepler’s laws for planetary systems. Intriguingly, the existence and properties of these dynamical laws were predicted in advance by MOND (Milgrom, 1983b), a paradigm in which the nonrelativistic laws of gravity and/or inertia are modified below a characteristic acceleration scale a_0 (see Famaey and McGaugh, 2012; Milgrom, 2014; Banik and Zhao, 2022, for reviews).

Interestingly, particle DM models and MOND theories make diametrically different predictions about how the properties of the dynamical laws (slope, normalization, and scatter) should depend on large-scale environment and cosmic time. In general, particle DM models predict that the dynamical laws should (1) have measurable intrinsic scatter because they emerge from the chaotic process of galaxy formation (Di Cintio and Lelli, 2016; Garaldi et al., 2018; Dutton et al., 2019), (2) evolve with cosmic time (Keller and Wadsley, 2017; Pillepich et al., 2019; Mayer et al., 2023), and (3) show “positive” residual correlations with large-scale environment (Paranjape and Sheth, 2022). Here, the term “environment” acquires the specific physical meaning of average gravitational acceleration due the large-scale structure in a given region of the Universe (Desmond et al., 2018). On the contrary, MOND predicts that the dynamical laws should (1) have zero or tiny intrinsic scatter (Milgrom, 1983b, 1994), (2) hold at any cosmic epoch (Hossenfelder and Mistele, 2018), and (3) show “negative” residual correlations with environment (Chae et al., 2020, 2021; Chae, 2022). To test these predictions, one needs high-quality rotation curves and mass models for statistical galaxy samples in different environments and at different redshifts. In Sect. 3, we described how the synergy

between HI surveys from SKA-Mid and optical/NIR images from state-of-the-art telescopes (e.g., Euclid, JWST, Rubin) will provide the necessary data to test these predictions.

2.2 The “cusp-core” and the “rotation-curve diversity” problems

According to DM-only cosmological simulations, CDM halos are singular (“cuspy”) at their centers with their volume density scaling as $\rho \propto r^{-1}$ (Dubinski and Carlberg, 1991; Warren et al., 1992), and can be described by a universal density profile, the NFW profile (Navarro et al., 1996). In contrast, the rotation curves of galaxies – especially those of DM-dominated low-mass and low-surface-brightness galaxies – often reveal much shallower density profiles or even constant density “cores” (de Blok, 2010; Katz et al., 2017; Read et al., 2017; Mancera Piña et al., 2025). This “cusp-core” discrepancy has been considered a major challenge to the standard CDM paradigm for a long time (see Bullock and Boylan-Kolchin, 2017, for a review). Proposed solutions either advocate for fundamental changes in the physics of the dark sector, or focus on the role of baryonic processes during galaxy formation and evolution. The first type of solutions comprises alternatives to CDM such as warm DM (e.g. Colín et al. 2000 but see also Macciò et al. 2012), self-interacting DM (e.g. Spergel and Steinhardt, 2000), and fuzzy DM made of ultralight axions (e.g. Goodman, 2000; Hu et al., 2000). The second type of solutions generally invokes strong baryonic feedback from supernova explosions and/or active galactic nuclei (AGN) to produce massive and repeated gas outflows, which lead to non-adiabatic oscillations in the gravitational potential, eventually transforming DM cusps into DM cores (e.g. Governato et al., 2010, 2012; Pontzen and Governato, 2012; El-Zant et al., 2016; Freundlich et al., 2020).

Some authors consider the cusp-core discrepancy “solved” in the Λ CDM context thanks to these ad-hoc adjustments in the modeling of the baryonic physics. However, different hydrodynamical simulations disagree on both the intensity of baryonic feedback processes and their effect on the distribution of DM, with some simulations predicting cores in a given mass range while others not forming cores at all (e.g., Bose et al., 2019). A general prediction of core-formation models via baryonic feedback is that the presence of a DM core would be correlated to the star formation history of a galaxy (so its baryonic-feedback history, e.g., Muni et al. 2025). This prediction could be tested by combining DM density profiles inferred from high-resolution HI rotation curves with, e.g., star formation histories inferred from observations at other wavelengths for a large sample of galaxies. On the other hand, self-interacting DM and fuzzy DM appear rather insensitive to baryonic processes (Robles et al., 2017) and predict different core-baryons relationships with respect to Λ CDM feedback models (Jiang et al., 2023), so there is hope of distinguishing between these two proposed solutions to the cusp-core problem.

A related issue in the Λ CDM context is the so-called “rotation-curve diversity” problem (Oman et al., 2015; Ghari et al., 2019). At a fixed galaxy mass, observations show a diversity in the inner rotation curve shapes – so in the DM density profiles – contrary to Λ CDM expectations if feedback processes were the same for galaxies of similar mass. More generally, the observed diversity of dwarf galaxies is not reproduced in current Λ CDM simulations, which may again have to do with baryonic feedback and the processes involved in DM core formation. For example, the NIHAO simulations produce both DM cores and a population of “ultra-diffuse” galaxies, but no compact dwarfs (Jiang et al., 2019). Importantly, the observed diversity in the inner rotation-curve shapes

of dwarf galaxies correspond to a diversity in their inner baryonic surface densities (see Lelli et al., 2014a), as quantified by the CDR (Fig.2, middle panel). The true underlying problem, therefore, is understanding the origin of the empirical baryon-DM coupling (Sect. 2.1). This issue was also pointed out by Sancisi (2004) more than 20 years ago: “why, in systems of comparable luminosity, are some DM halos cuspy (like the light) and others (also like the light) are not?”.

Another related issue in the Λ CDM context is the “too-big-to-fail” problem, whereby the inner DM density of satellite galaxies is lower than expected in cosmological simulations, suggesting that some of the most massive DM sub-halos did not form stars as much as they should have (e.g. Sales et al., 2022). To a large extent, this problem is a combination of the “cusp-core” and the “missing satellites” problems, extrapolated to the lowest mass galaxies (Bullock and Boylan-Kolchin, 2017). Obtaining HI rotation curves for large samples of gas-rich satellite galaxies may further clarify this problem and its relation to the “cusp-core” and “rotation-curve diversity” problems.

Finally, we point out that these various problems (“cusp-core”, “rotation curve diversity”, “too-big-to-fail”) are specific to Λ CDM and do not apply in a modified gravity context. In general, modified gravity theories are required to reproduce the observed rotation curves starting from the observed baryonic distribution, so the diversity in the rotation curve shapes (e.g., steeply or slowly rising) naturally corresponds to a diversity in the baryonic mass distribution (e.g., cuspy or cored). Yet, the inner parts of rotation curves remain very interesting in the modified-gravity context because it is the regime in which different theories make somewhat different predictions (e.g., Lelli et al., 2017; Petersen and Lelli, 2020; Chae, 2022). Constructing a large sample of galaxies with high-resolution HI data (providing both the rotation curve and the gas distribution) and ancillary optical/NIR photometry is crucial to distinguish between different modified-gravity theories.

3 Observational prospects with SKA-Mid

HI surveys with the SKA-AA4 will allow us to address the open questions described in the previous section. Different questions, however, require different survey strategies. For example, to study the inner shapes of rotation curves in high detail, we need targeted HI observations of nearby galaxies at resolutions of $1''$ - $2''$ and 5 km/s reaching HI column densities of 5 - 10 $M_{\odot}$ pc^{-2} (typical of the innermost galaxy regions). Instead, to investigate the possible dependency of the dynamical laws and DM properties on cosmic environment, we need a large statistical sample (tens of thousands of galaxies) with “standard” medium-shallow HI data, e.g., reaching a 3σ HI column density of 1 $M_{\odot}$ pc^{-2} at “typical” resolutions of $10''$ - $20''$ and 5 – 10 km/s. Finally, to investigate the possible evolution of the dynamical scaling laws with z , we need an ultra-deep pencil-beam survey in a well-studied cosmological field. This section describes these potential HI surveys and highlights the importance of the SKA-Mid AA4 to reach the requested resolutions and/or sensitivities.

3.1 Reaching higher spatial resolutions

For many years, the main limitation of HI rotation curves has been the relatively low angular resolution of $15''$ – $30''$ achievable with “historical” radio interferometers, such as the Westerbork Synthesis Radio Telescope (WSRT), the Very Large Array (VLA), and the Australia Telescope Compact Array (ATCA). Such angular resolutions are significantly worse than those of $1''$ – $2''$

available for the complementary optical/NIR images (used to compute the stellar gravitational contribution) and/or optical spectroscopy (used to measure $H\alpha$ rotation curves). Having kinematic data at high angular resolution is crucial to trace accurate rotation curves in the inner parts, minimizing the so-called beam-smearing effects (see Sect. 4.2.1) as well as detecting and modeling potential noncircular motions (see Sect. 4.3.2). Indeed, understanding these two observational effects has been important in establishing the presence of DM cores in low-mass galaxies (see, e.g., de Blok, 2010), contrary to the original Λ CDM prediction of DM cusps (see Sect., 2.2).

In this context, a major advance was represented by the THINGS survey (Walter et al., 2008), which used the Very Large Array (VLA) in multiple configurations (B, C, D) to reach angular resolutions of $\sim 6''$ in 34 nearby galaxies (albeit rotation curves could be derived for only 19 of them, see de Blok et al., 2008). After THINGS, several other HI surveys used the VLA with a similar strategy (e.g., Hunter et al., 2012; Ott et al., 2012; Lelli et al., 2014a). To our knowledge, to date, the highest angular resolution achieved with HI observations is $2''$ using VLA data in A, B, C, D configurations (Lelli et al., 2012). These types of high-resolution HI observations have been quite successful but require long observing campaigns spanning over several years, so that the VLA antennas are moved into different array configurations.

Thanks to its long baselines, SKA-Mid AA4 will completely change the game because it will allow us to routinely obtain HI data at $1'' - 2''$ resolution. For example, at the higher redshift limit ($z \simeq 0.2$) of AA4's band 2, we will reach HI column density sensitivities of $\sim 1 \text{ M}_\odot \text{ pc}^{-2}$ at spatial and spectral resolutions of $1'' - 2''$ and $\sim 5 \text{ km s}^{-1}$, respectively, in under 15 hours on-source observing time. At the lower redshift end ($z \simeq 0$), the same type of data will be achieved in just ~ 10 hours. For example, it will be conceivable to perform a legacy survey of 20-30 nearby galaxies in about 200-300 hours of SKA-Mid to obtain HI data with similar angular resolutions of $H\alpha$ integral-field spectroscopy (e.g., from MaNGA or SAMI), CO interferometry (e.g., from ALMA), and NIR imaging (e.g., from Spitzer). This type of high-resolution HI survey may have plenty of scientific goals, such as relating the HI distribution and kinematics with that of molecular and ionized gas to understand the star-formation process. In terms of the DM problem, there will be two major advances. Firstly, we could measure the properties of DM cores (sizes, volume densities) with unprecedented accuracy, even in the smallest HI-bearing galaxies, providing major constraints to Λ CDM models of galaxy formation as well as to alternative DM models (e.g., warm DM, self-interacting DM, fuzzy DM). Secondly, we could test different modified-gravity theories in the regime where they give the largest differences (e.g., Petersen and Lelli, 2020; Chae, 2022).

3.2 Building larger galaxy samples

A comprehensive understanding of the relationship between cosmic environment and galaxy dynamics (rotation-curve shapes, DM content, and so on) requires large, homogeneous, statistical samples of spatially resolved galaxies across a full range of environmental densities, from rich clusters and loose groups to truly isolated field galaxies. As we mentioned in Sect. 1, existing HI samples of spatially resolved galaxies have limited statistics and do not capture the full environmental diversity necessary to study the possible relations between the DM content of galaxies and their location in the large-scale structure of the Universe.

With the SKA-Mid AA4, it will be possible to spatially resolve the HI kinematics of large galaxy samples in a time-efficient way. To this aim, there are two possible survey strategies:

1. A “blind” HI survey of a large fraction of the sky, similarly to APERTIF (Adams et al., 2022) and WALLABY (Westmeier et al., 2022), but reaching higher angular resolutions ($\sim 5''$ versus $\sim 15'' - 30''$) at similar 5σ column density sensitivities ($\sim 1 \text{ M}_\odot \text{ pc}^{-2}$ over 20 km s^{-1}), so that a much larger number of HI disks will be spatially resolved;
2. A “snapshot” targeted survey of $\sim 20,000$ galaxies with known optical/NIR properties (e.g., galaxies with Euclid images), so that the integration time can be individually set according to the desired sensitivity and angular resolution to resolve the HI disk. The target sample of 20,000 galaxies is two orders of magnitudes larger than SPARC and about one order of magnitude larger than what we expect from WALLABY, APERTIF, and BIG-SPARC.

Given that SKA-Mid will have standard single-pixel receivers, the field of view (FoV) at 21 cm remains rather limited, so it is not immediately clear which of the two survey strategies is more effective, also taking into account commensal science goals.

In the case of a “blind” wide-field medium-shallow HI survey with the Band 2 of SKA-Mid AA4, we can make a rough estimate considering the HI mass function of Jones et al. (2018) and the well-studied HI mass-size relation (e.g., Wang et al., 2016). If we set a 3σ column density sensitivity of $1 \text{ M}_\odot \text{ pc}^{-2}$ for a synthesized beam of $4.3''$ (Briggs’ robust of 1) and spectral channels of 3.1 km s^{-1} , we expect that we will spatially resolve about 20,000 galaxies by covering 150 square degrees in about 900 hours of SKA-Mid AA4. Importantly, a spatial resolution of $4.3''$ will allow resolving galaxies with a HI mass of $\sim 10^7 \text{ M}_\odot$, pushing down in terms of the BTFR and other dynamical scaling laws. However, the drawback is that only a few low-mass galaxies will be resolved by such a “blind” HI survey, while the sample will be heavily dominated by galaxies with HI masses near the knee of the Schechter function. Thus, the dynamical scaling laws will not be sampled in a uniform way across their dynamic ranges, limiting our accuracy in determining their slope and/or intrinsic scatter as a function of mass, rotation velocity, surface density, and so on.

A pointed HI survey may be more time-expensive than a wide-field HI survey to spatially resolved 20,000 galaxies. For example, if we consider a mean integration of 15 minutes per galaxy, it will take 5000 hours to observe 20,000 galaxies. This time estimate is quite conservative: the exact numbers will depend on the exact definition of the galaxy sample and may be substantially smaller. Indeed, a single SKA-Mid pointing may simultaneously probe multiple galaxies in clusters and/or groups, maximizing the survey efficiency. A lower survey efficiency may occur only for targeting truly isolated galaxies in voids, albeit several background and foreground galaxies may be detected and spatially resolved along the line of sight. A pointed HI survey will have the advantage of building a carefully selected galaxy sample that *uniformly* covers a broad range in baryonic mass, surface brightness, star-formation rates, environmental properties, and so on. In addition, a pointed HI survey may maximize the synergy with other existing multi-wavelength data, such as (1) imaging from HST, Spitzer, JWST, Euclid, Rubin, and so on, (2) availability of precise distances from Cepheids, the tip of the red giant branch method, supernovae Ia, surface brightness fluctuations, and so on, and (3) availability of complementary kinematic data from H α integral-field

spectroscopy and/or interferometric CO observations.

Thinking farther ahead in the future, installing phase-array feeds (PAF) on SKA-Mid antennas (similarly to ASKAP and APERTIF) will substantially enlarge the instantaneous FoV of the telescope. In this way, wide-field “blind” surveys will be the obvious operational mode to increase galaxy samples with spatially resolved HI kinematics by another 1-2 orders of magnitudes.

3.3 Moving towards higher redshifts

The intrinsic faintness of the HI emission line, combined with the limited spectral bandpasses and persistent radio frequency interference in the L-band, have historically limited HI detections of individual galaxies to the local universe ($z \simeq 0$). This paradigm is beginning to shift with a new generation of HI surveys with SKA pathfinders (Ponomareva et al., 2021; Xi et al., 2024; Smith, 2024; Jarvis et al., 2025), but the existing facilities do not have the requested sensitivity to map galaxies beyond $z \gtrsim 0.1$ with enough angular resolution to enable kinematic modeling.

At intermediate redshifts ($z \simeq 0.3 - 1$) and cosmic noon ($z \simeq 1 - 3$), rotation-curve studies currently rely on different tracers, such as ionized gas using integral-field spectroscopy of the H α , [O II], and [O III] lines (e.g., Di Teodoro et al., 2016, 2018; Price et al., 2021; Puglisi et al., 2023), or molecular gas using mm/submm interferometric observations of CO and [C I] lines (e.g., Lelli et al., 2023; Rizzo et al., 2023; Lin et al., 2025). Integral-field observations with the ESO Very Large Array (VLT) are relatively time efficient, so they allowed resolving the H α and [O II] kinematics in several hundreds of galaxies (e.g., Tiley et al., 2019; Sharma et al., 2025), whereas mm/submm interferometry with ALMA is more time costly, so galaxy samples with resolved CO or [C I] kinematics remains limited to a few tens of objects. On the other hand, ionized gas probes an intrinsically hotter ISM phase than molecular or atomic gas, so the pressure support to the gravitational potential becomes important to properly model (see Sect. 4.1.2). Moreover, non-circular motions (inflows, outflows) may be affecting the ionized gas kinematics in more complex ways than for the neutral gas (e.g., Lelli et al., 2018). Most importantly, both ionized and molecular gas tracers have an intrinsic physical limitation: the ionized and molecular gas disks of galaxies generally do not extend beyond their stellar components, so one can only trace rotation curves with a limited radial extent (see Fig. 3).

In relatively massive ($M_\star > 10^{10} M_\odot$), the inner high-acceleration regions are generally baryon dominated, so it is not possible to firmly probe the DM effect in cosmic-noon galaxies due to the disk-halo degeneracy: given the significant uncertainties on stellar and gas masses, mass models with only baryons and mass models with a DM halo can fit the observed rotation curves with similar accuracy (Lelli et al., 2023). The current observational situation at cosmic noon is similar to that at $z \simeq 0$ in the 1970’s-1980’s (e.g., Kalnajs, 1983; Kent, 1986): one needs extended HI rotation curves to firmly probe the DM effect in the outer parts (van Albada et al., 1985; Kent, 1987).

In lower mass galaxies ($M_\star < 10^{10} M_\odot$), the inner rotation curves may potentially start probing the DM effect, similarly to DM-dominated low-mass galaxies at $z \simeq 0$ (see Lelli, 2022, for a review). However, molecular lines (CO and [C I]) are hard to detect in such low-mass galaxies, even with ALMA. Regarding ionized gas kinematics, the typical spatial resolution of integral-field spectrographs (0.5 – 1 arcsec) becomes a major limitation in low-mass galaxies, together with the limited spectral resolution of $R = \lambda/\Delta\lambda \simeq 3500 - 4000$, which corresponds to $\sigma \simeq 40$ km/s for

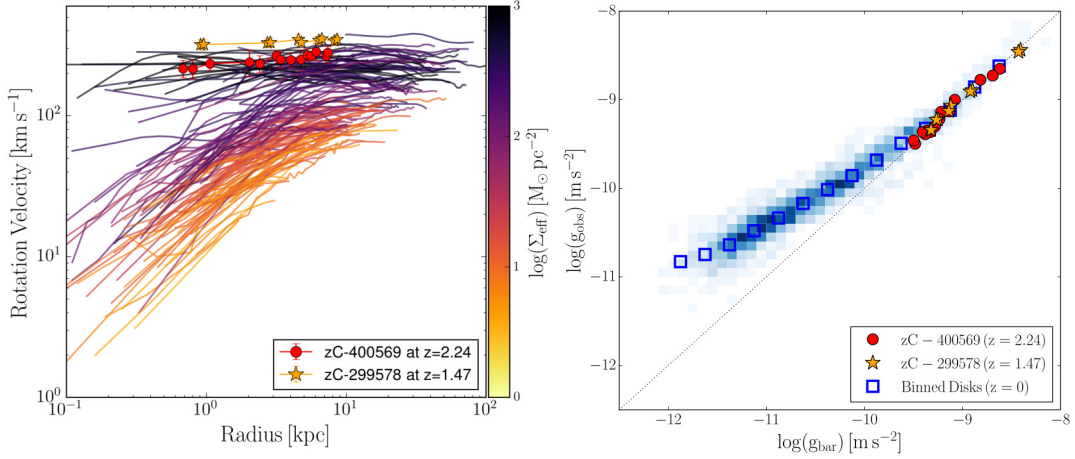


Figure 3: Comparison between the dynamical properties of galaxies at $z \simeq 0$ from the SPARC database with two galaxies at cosmic noon with top-quality VLT/SINFONI and ALMA data (adapted from Lelli et al., 2023). **Left panel:** The rotation curves of cosmic-noon galaxies are much less extended than those of local galaxies of comparable mass and rotation velocity due to the lack of HI observations. **Right panel:** The location of cosmic-noon galaxies on the radial acceleration relation at $z \simeq 0$ shows that their H α +CO rotation curves are probing only the inner, high-acceleration, baryon-dominated regions. To firmly probe the DM effect, we need extended HI rotation curves probing the outer, low-acceleration, DM-dominated regions.

the H α line at $z \approx 1$ (see, e.g., Stott et al., 2016; Wisnioski et al., 2019). All these facts preclude the study of the dynamical scaling laws at low densities and low accelerations, which provide the strongest tests of Λ CDM models against alternative theories.

SKA-Mid AA4 will revolutionize our view of the HI kinematics in galaxies at $z > 0.1$ thanks to its unique combination of angular resolution and sensitivity. To trace a reliable HI rotation curve out to the outer DM-dominated regime, a column density sensitivity of $\sim 1 \text{ M}_\odot/\text{pc}^2$ is required. In addition, to constrain the geometric parameters of the HI disk (see Section 4), the disk major axis should be sampled with at least 5 resolution elements. If we (1) adopt this combination of sensitivity and resolution as a metric for measuring HI rotation curves, (2) assume the local HI mass-size relation (e.g., Wang et al., 2016) to predict the size of the HI disk, and (3) account for cosmological redshift broadening and dimming (Meyer et al., 2017), we find that a deep SKA-Mid AA4 integration is capable of mapping Milky Way-mass galaxies ($M_{\text{HI}} \simeq 10^{10} \text{ M}_\odot$) out to $z \simeq 1$, when the Universe was only half its age.

Figure 4 illustrates the dramatic leap in our ability to push kinematic studies of HI disks out in redshift with a deep ($\sim 10,000$ hour) SKA-AA4 Band 1 pencil beam survey, which reaches the required sensitivity and resolution at $z \sim 1$ to map massive HI disks. In combination with a Band 2 pencil beam SKA-AA* survey to map disks across the redshift range $0 \lesssim z \lesssim 0.45$, the SKA would enable dynamical studies of HI disks across cosmic time, for the first time. The SKA is the only facility under construction or being planned that will have this capability.

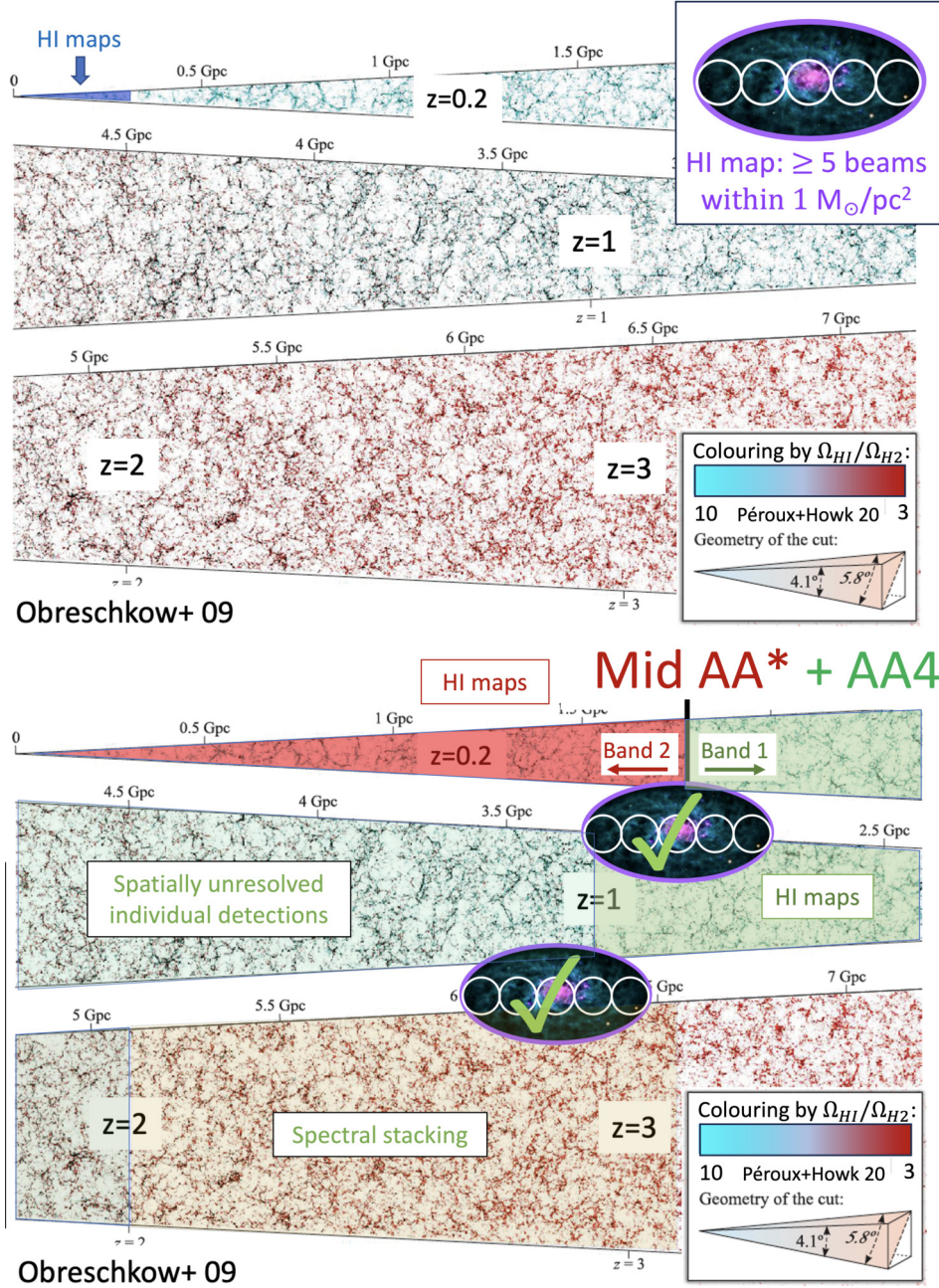


Figure 4: Illustration of the state of the art in detecting HI emission in individual galaxies across cosmic time today (top panel), and with an SKA Key Science Project (KSP; bottom panel). In each panel, a simulated distance wedge of the Universe from Obreschkow et al. (2009) is shown, with individual galaxies represented by dots with colours that approximate the universal ratio of HI to H_2 at that cosmic time (Péroux and Howk, 2020). Our census of HI in emission in individual galaxies today is limited to maps in the local universe ($z \lesssim 0.1$). A $\sim 10,000$ -hours pencil-beam KSP survey with SKA Mid AA4 could map HI disks with HI masses $\log(M_{\text{HI}}/M_{\odot}) \gtrsim 10$ out to $z \approx 1$, while spatially unresolved individual detections would extend to $z \approx 2$, and stacks to the edge of Band 1 at $z \approx 3$.

4 Modeling prospects with next-generation software

4.1 State-of-the-art on kinematic fitting tools

The H I component of galaxies can be described by a thin rotating disk with nearly circular orbits. In this approximation, the projected line-of-sight velocity $V_{\text{l.o.s.}}$ at a sky position (x, y) is given by

$$V_{\text{l.o.s.}}(x, y) = V_{\text{sys}} + V_{\text{rot}}(R) \sin(i) \cos(\theta), \quad (1)$$

where V_{sys} is the galaxy systemic velocity (due to the Hubble expansion and peculiar motions), V_{rot} is the rotational velocity at radius R in the disk plane, i is the inclination angle between the normal to the disk plane and the line of sight, and θ is the azimuthal angle in the disk plane given by

$$\cos(\theta) = \frac{-(x - x_0) \sin(P.A.) + (y - y_0) \cos(P.A.)}{R}, \quad (2)$$

where (x_0, y_0) are the sky coordinates of the disk center and $P.A.$ is the position angle taken in anti-clockwise direction between the North direction and the receding half of the major axis of the projected disk. Thus, the rotation curve $V_{\text{rot}}(R)$ can be measured knowing six geometric parameters: V_{sys} , x_0 , y_0 , $P.A.$, and i . In the following, we describe standard methods to measure $V_{\text{rot}}(R)$ from emission-line data: tilted-ring and parametric modeling implemented via two dimensional (2D) or three dimensional (3D) approaches.

4.1.1 Basic techniques: tilted-ring and parametric modeling

The tilted-ring modeling approach (Rogstad et al., 1974; Begeman, 1989) divides the gas disk into a number of rings (usually set by the spatial resolution of observations). In principle, the parameters in Eq. 1 and 2 can be fitted independently in each ring. However, this ring-by-ring fit is often under-determined and/or degenerate, so the best-fit parameters may unrealistically jump from ring to ring. One approach is that, after a first preliminary fit, (x_0, y_0) and V_{sys} are measured as the mean (or median) values across the rings and kept fixed in subsequent iterations. The angles $P.A.$ and i can be estimated in the same way unless the gas disk is warped: in such cases, $P.A.$ and i vary from ring to ring and their radial trend can be appropriately modeled with a smooth function to obtain a reliable $V_{\text{rot}}(R)$ without unrealistic jumps. Another approach is to perform a global fit to the disk by grouping various parameters (such as x_0 , y_0 , and V_{sys}) to have a single value across all rings. The details of these implementations depend on the specific tilted-ring fitting software.

The parametric modeling approach uses parametric functions to describe $V_{\text{rot}}(R)$, such as an arctan function (Epinat et al., 2010), and fits a global model to the whole disk. In several circumstances, parametric modeling can require fewer free parameters to describe $V_{\text{rot}}(R)$ than typical tilted-ring models. This eases the use of Markov Chain Monte Carlo (MCMC) algorithms (Bouché et al., 2015; Davis et al., 2017b) and neural networks (Dawson et al., 2021) in a Bayesian framework (see Sect. 4.4.1 and 4.4.2). On the other hand, the parametric modeling cannot account for disk warps in an automated way because specific functions for the radial variation in $P.A.$ and i need to be defined a-priori. Parametric modeling may also neglect possible real features in the rotation curve because a smooth shape is imposed on $V_{\text{rot}}(R)$ a-priori. While tilted-ring models provide the most empirical derivation of the rotation curve itself, some modeling tools allow for hybrid approaches, mixing tilted-ring-style approaches for some parameters and parametric functions for others.

Special types of parametric models are those in which V_{rot} is described by constraining the gravitational potential of different components (stars, gas, and DM halo), so the fitting involves physical parameters (such as masses and characteristic radii) together with geometric ones. Ideally, this approach permits to build dynamically self-consistent models (for instance, by assuming hydrostatic equilibrium) because the properties of the gravitational potential and those of the gas tracer are constrained at the same time. In addition, this approach bypasses issues in the mass modeling of rotation curves due to correlations between $V_{\text{rot}}(R)$ measurements (Posti, 2022). The main drawback of this approach is its reliance on pre-implemented (cosmologically-motivated) models for the DM distribution, which may not capture the observational reality and/or the full complexity of galaxy dynamics. In addition, these parameteric models have relatively heavy computational costs when dynamical self-consistency is requested, and they require refitting the full data (in 2D or 3D) several times for testing different DM halos or modified-gravity theory. On the contrary, it is computationally fast and conceptually straightforward to fit different DM halos and/or modified gravity theories to the empirical 1D rotation curve from a tilted-ring model. Although promising, HI modeling tools in a dynamically self-consistent framework have not yet been implemented.

4.1.2 From rotation velocity to circular velocity

When modeling the mass distribution of galaxies, we are interested in the circular velocity of a test particle ($V_c^2 = -R \nabla \Phi$) tracing the total gravitational potential Φ . In general, this is given by a combination of rotation (V_{rot}) and pressure ($\rho \sigma_V^2$) support. Contrary to the case of stellar kinematics and/or ionized gas, pressure support is often negligible in HI disks of spiral galaxies because the velocity dispersion is typically between $5 - 15 \text{ km s}^{-1}$, so effectively $V_c \simeq V_{\text{rot}}$.

For galaxies in which $V_{\text{rot}}/\sigma_V \simeq 1 - 4$ (e.g., dwarf galaxies), pressure support can be taken into account using the so-called asymmetric-drift correction. Assuming that (i) Φ is axisymmetric depending only on the cylindrical coordinates (R, z) , and (ii) the velocity dispersion σ_V is isotropic as expected for collisional gas, the circular velocity is given by (Binney and Tremaine, 1994)

$$V_c = \sqrt{V_{\text{rot}}^2 - \sigma_V^2 \left(\frac{\partial \ln \rho}{\partial \ln R} + \frac{\partial \ln \sigma_V^2}{\partial \ln R} \right)}, \quad (3)$$

where $\rho(R, z)$ is the gas volume density, which is not directly observable.

A common approach (Meurer et al., 1996; Lelli et al., 2014b) is to assume that the vertical density distribution does not depend on radius, so $\partial \ln \rho / \partial \ln R = \partial \ln \Sigma / \partial \ln R$, where the gas surface density $\Sigma(R)$ can be traced from emission-line maps. Then, one may either fit the product $\Sigma \sigma_V^2$ with a smooth functional form (Iorio et al., 2017) or treat $\Sigma(R)$ and $\sigma_V(R)$ independently. The main caveat with this approach is that the thickness of the gas disk is expected to increase with radius from theoretical considerations based on vertical hydrostatic equilibrium (Romeo, 1992; Olling, 1995; Bacchini et al., 2020).

Another approach (Burkert et al., 2010) is to use the Spitzer's solution for a self-gravitating disk in vertical equilibrium (Spitzer, 1942), so $\rho(R, z) = \rho_0(R) \text{sech}^2(z/z_d)$ where $\rho_0(R) = \pi G \Sigma(R)^2 / 2 \sigma_V(R)$ is the gas volume density at $z = 0$. The gaseous disks of galaxies, however, are not self-gravitating. A self-consistent asymmetric-drift correction should use an iterative

approach to (i) obtain $\Phi(R, z)$ fitting $V_c(R)$ with a mass model for a given DM halo shape, and (ii) measure the variation of the disk thickness with radius by solving the vertical hydrostatic equation for the given $\Phi(R, z)$ (Bacchini et al., 2020). This approach may be implemented in future tools and initial efforts are underway (e.g., Mancera Piña et al., 2025). It may also be possible to calibrate the asymmetric-drift correction using numerical simulations, for example to obtain prescriptions informed by galaxy morphology (Kretschmer et al., 2021). Parametric models built in the dynamically self-consistent framework can potentially account for the flaring of the disk, providing direct computation of pressure support via Eq. 3.

At any rate, the dominant uncertainty in the asymmetric-drift correction is driven by $\sigma_v(R)$, not by the vertical density distribution. In fact, one often has to assume that σ_v is constant with radius because the velocity dispersion profile cannot be accurately measured in the lowest mass galaxies. Then, for an exponential gaseous disk with scale length R_d , Eq. 3 simplifies to

$$V_c = \sqrt{V_{\text{rot}}^2 + \sigma_v^2(R/R_d)}. \quad (4)$$

Eq. 4 can be used as a rough first-order approximation in poorly resolved galaxies, which are the focus of the next section.

4.2 Modeling of poorly resolved disks

HI surveys with SKA-Mid may provide tens or even hundreds of thousands of galaxies that are just barely resolved. Unsurprisingly, preliminary results from the WALLABY survey indicate that barely resolved systems will dominate any wide-field HI survey: only $\sim 10\%$ of the HI detections are sufficiently resolved for detailed kinematic modeling (Westmeier et al., 2022; Murugesan et al., 2024). HI modeling in the barely resolved regime demands specific techniques, such as the full treatment of beam smearing in 3D (Sec 4.2.1), as well as algorithms for the automated acceptance/rejection of kinematic fits (Sec. 4.2.2).

4.2.1 Beam-smearing effects: 2D and 3D approaches

Tilted-ring and parametric modeling can be applied either to 2D velocity maps ($V_{l.o.s.}$ at each sky position) or directly to the 3D emission-line cubes. The 2D approach is robust for well-resolved galaxies with tens or hundreds of resolution elements (Begeman, 1989), but becomes unreliable in less resolved galaxies as is often the case for dwarf galaxies (Swaters et al., 2009). In this situation, the derivation of velocity maps is ambiguous due to beam-smearing effects, altering the intrinsic shape of the line profiles. Gas emission with different $V_{l.o.s.}$ is flux-averaged within the resolution element, so the line profiles tend to become broad and non-Gaussian with long tails of emission towards the galaxy systemic velocity. In such cases, if one estimates $V_{l.o.s.}$ as the moment-one along the profile or as the central velocity of a Gaussian fit, the rotational velocity will be underestimated, especially in the central parts (Swaters et al., 2009; Lelli et al., 2010). Moreover, if one estimates the gas velocity dispersion σ_v as the moment-two along the profile or as the width of a Gaussian fit, the intrinsic velocity dispersion will be severely overestimated (Di Teodoro and Fraternali, 2015). For galaxies with high inclinations, the minor axis is always poorly resolved due to the disk projection along the line of sight. Therefore, similar effects occur when analyzing edge-on galaxies, limiting 2D methods to galaxies with inclinations $< 75^\circ$. There is no simple way to fully correct for beam

smearing using a 2D approach because it depends on the interplay between intrinsic rotation-curve shape, gas distribution, disk thickness, and disk inclination (Begeman, 1989).

The 3D approach builds model cubes using Eq. 1 together with the gas surface density profile, the disk thickness, and the intrinsic σ_V (Corbelli and Schneider, 1997; Sicking, 1997). The model cubes are then smoothed to the same spectral and spatial resolution of the observations (reproducing beam smearing effects) and compared with the observed one until a best-fit is found. The 3D approach can then recover the intrinsic $V_{\text{rot}}(R)$ and $\sigma_V(R)$ profiles (Józsa et al., 2007; Swaters et al., 2009; Lelli et al., 2010; Di Teodoro and Fraternali, 2015). Over the past years, several 3D parametric fitting tools have been developed, such as KinMS (Davis et al., 2013), GalPak-3D (Bouché et al., 2015), DysmalPy (Price et al., 2021; Lee et al., 2025) and Geko (Danhaive and Tacchella, 2025). Importantly, the 3D parameteric modelling may not capture the correct beam-smearing effects in poorly resolved galaxies because it imposes a specific rotation-curve shape. For example, an intrinsically flat rotation curve leads to more severe beam-smearing effects than an intrinsically solid-body one. A 3D tilted-ring modeling is often the best solution, and several titled-ring 3D fitting tools are now available, such as TiRiFiC (Józsa et al., 2007) and its automated version FAT (Kamphuis et al., 2015), ^{3D}Barolo (Di Teodoro and Fraternali, 2015), and a new version of KinMS (Davis et al., 2020). These 3D kinematic modeling tools allows for the derivation of rotation curves even for poorly resolved galaxies (with only $\sim 3 - 4$ along the major axis), which cannot be reliably modeled in 2D. However, even these packages may still be affected by beam-smearing, making the innermost rings still somewhat uncertain (see for example Deg et al. 2022). In principle, the best way to remove many of the uncertainties discussed above would be to model the emission of the kinematic tracer in the uv -plane, as discussed in Section 4.4.3.

4.2.2 Automated acceptance/rejection criteria

A key use of untargetted surveys is the comparison with large numerical simulations. A significant challenge in these comparisons is understanding the “selection function”. That is, how to select simulated galaxies that *could* have been observed and modeled in the same way as a given survey. Such comparisons require that the same modeling approach and software is applied to both mock and real observations. The traditional approach of accepted models based on visual inspection (e.g., Deg et al., 2022) introduces an hard-to-quantify effect to the selection function. Moreover, even if this effect could be quantified, the large galaxy samples expected from SKA-Mid surveys imply that visual inspections will become unfeasible. However, visual inspection by expert users may still play an important role: for example, one may train a machine-learning algorithm using the assessments from visual inspection as a robust training set.

An automated approach based on a suite of test models/observations will need to be adopted for wide-field HI surveys, where the majority of detections are marginally resolved. This approach must take into account the limits of modeling software in terms of resolution, inclination, uncertainties, and even galaxy mass or rotation velocity. The reliability of kinematic models strongly depends on galaxy mass and inclination, being progressively worse for low-inclination, low-mass galaxies. For example, Deg et al. (2025) demonstrated that, at velocities below $\sim 50 \text{ km s}^{-1}$, kinematic models for galaxies resolved by 5 beams were reliable only for inclinations larger than 40° (see also Rhee et al., 2004; Kuzio de Naray et al., 2006; Read et al., 2016).

One challenge in any automated approach is the presence of “contaminants” in the final catalog, i.e., rotation curves that do not represent the underlying gravitational potential of the galaxy. These contaminants can be found both in catalogs generated by visual inspection (see for example some of the low-velocity WALLABY models of [Deg et al. 2022](#)) and in automated approaches. Automated accept/reject rules can be tuned to increase the number of total models at the cost of an increased contaminant fraction, so one needs to find the best compromise to minimize the contaminant fraction while maximizing the total number of acceptable models. Despite the numerous challenges of modelling large numbers of marginally resolved, poor S/N detections, the scientific benefits are well worth the effort. They will provide statistically significant samples of kinematic models across a wide range of environments and, in the SKA era, a range of redshifts.

4.3 Modeling of highly resolved disks

SKA-Mid AA4 will allow for HI observations at $1''$ - $2''$ resolution. These high-resolution data will probably uncover details in the gas distribution and kinematics that need to be modeled with new approaches. These may include highly asymmetric gas distributions (clouds, clumps, shells, holes, and so on), complex non-circular motions (inflows, outflows, and so on), and a complex vertical structures (extraplanar gas and disk flares). We discuss potential ideas to tackle these challenges.

4.3.1 Asymmetric gas distributions

Where gas distributions are clumpy and/or strongly asymmetric, traditional tilted ring methods (which impose axisymmetry) will fail to reproduce the observed gas distribution. The impact of this effect on the resulting rotation curve can vary wildly. In systems that are clumpy, but where emission fills much of the radial ring, axisymmetric model generally perform well (e.g., [Davis, 2014](#)). However, in extreme cases, the best-fit model may create artificial spikes in the rotation curve to lower the surface brightness in a given ring, a clearly pathological behavior. Various techniques can be used to avoid such issues, but not without some drawbacks.

Three main techniques have been used to deal with clumpy distributions. The simplest one is to normalize the flux in a given spatial pixel of the 3D model to match the observed intensity (moment-zero) map a-posteriori ([Lelli et al., 2012](#)). The drawback of this method is that it does not fully account for beam-smearing effects from the intrinsic variations in the surface density distribution within the beam. The second technique mitigate this effect by using the “clean components” determined during the imaging of the uv data to determine the flux distribution, and fitting the kinematics of these components ([Smith et al., 2019](#)). The “trusted” flux in the model is then reproduced by definition. The third technique adopts a more drastic solution, that is allowing the kinematics and positions of a cloud of point sources to be fitted to the data. While very powerful, this technique could have millions of free parameters and is thus extremely computationally expensive. All these methods, however, removes nearly all the constraining power on the disk inclination because they are forced to reproduce the observed gas distribution. The best strategy, therefore, is a two-step approach: first fitting a model that is axisymmetric in both gas distribution and kinematics to constrain the disk geometry (especially the inclination), then fitting a model that is not axisymmetric in the gas distribution to constrain the rotation curve ([Lelli et al., 2012](#)).

In the future, new machine learning-based techniques such as diffusion models may be used to

determine the gas distribution, although such efforts are in their early stages at the present time. Another way to account for clumpy gas distribution may be to do kinematic modeling directly in the uv plane, modeling all the recoverable scales at the same time, as we discuss in Sect. 4.4.3.

4.3.2 Non-circular motions in the disk plane

The simplest form of non-circular motion is a radial flow (V_{rad}) in the disk plane, giving an additional term $V_{\text{rad}}(R) \sin(i) \sin(\theta)$ in the right part of Eq. 1. More generally, V_{rot} and V_{rad} are the first-order terms of an harmonic expansion of the line-of-sight velocity (Schoenmakers et al., 1997):

$$V_{\text{l.o.s.}}(x, y) = V_{\text{sys}} + \sum_{m=1} [c_m(R) \cos(m\theta) + s_m(R) \sin(m\theta)], \quad (5)$$

where the harmonic order $m = 1$ gives $c_1 = V_{\text{rot}} \sin(i)$ and $s_1 = V_{\text{rad}} \sin(i)$. A perturbation of the gravitational potential of order m causes harmonics of order $m - 1$ and $m + 1$ in projection (Schoenmakers et al., 1997). For example, bar-like and oval distortions with $m = 2$ give $m = 1$ and $m = 3$ terms in $V_{\text{l.o.s.}}$ (Spekkens and Sellwood, 2007). The harmonic decomposition increases the number of free parameters in tilted-ring fits, so it can only be applied to high-quality observations. This method has been mostly applied in 2D (Gentile et al., 2005, 2007; Trachternach et al., 2008; Marasco et al., 2018), but further progress can be made with a 3D approach (Hogarth et al., 2024). In particular, it will be important to investigate strategies to tackle the various degeneracies between the free parameters of the harmonic expansion, for example by using a Bayesian approach with physically motivated priors, as those discussed in Sect. 4.4.1.

4.3.3 Vertical structure: extraplanar gas and flares

A non-negligible fraction ($\simeq 15\%$) of HI in star-forming galaxies resides in a thick structure that surrounds the thin disk, extending typically a few kpc above and below the midplane. This “extraplanar” HI emerged by performing deep HI observations, reaching column densities of $0.1 \text{ M}_{\odot} \text{ pc}^{-2}$ or less (Fraternali et al., 2002; Oosterloo et al., 2007), which will be routinely reached by SKA-Mid. The extraplanar HI gas is often characterized by peculiar kinematics, showing a slower rotation than the HI in the thin disc and a global inflow motion with typical speeds of a few tens km s^{-1} (e.g. Fraternali et al., 2002; Oosterloo et al., 2007; Marasco et al., 2019). These peculiar properties are broadly consistent with an origin from galactic fountain triggered by stellar feedback (Fraternali, 2017). To date, the peculiar kinematics of the extraplanar HI gas has been modeled in 3D with substantial human intervention, given the increased number of free parameters and overall complexity. In the SKA era, we may need to explore automated approaches to account for extraplanar HI gas in large galaxy samples.

As discussed in Sect. (4.1.2), HI discs are naturally flared, with a scale-height that is expected to reach up to ~ 1 kpc in the outskirts (Bacchini et al., 2020; Mancera Piña et al., 2022). High-resolution observations should be able to distinguish between extraplanar gas and flaring of the HI disc because these features occur in distinct regions (inner versus outer disk, respectively) and have different kinematics. In fact, extraplanar gas is often visible as residual emission after the subtraction of the best-fit kinematic model, including potential disk flaring (e.g., Oosterloo et al., 2007). However, in marginally resolved galaxies, the contamination from extraplanar gas might be

inevitable, leading to a systematic underestimate of $V_{\text{rot}}(R)$. Although there is no obvious way to circumvent this issue, tests using mock degraded data of nearby galaxies will be valuable to quantify the magnitude of this effect and to elaborate effective correction strategies.

4.4 Next-generation modeling techniques

4.4.1 Bayesian techniques: improvements and optimization

Both parametric and non-parametric fitting methods make use of a substantial number of free parameters, which are often degenerate. The most important degeneracy is between V_{rot} and inclination, but additional degeneracies arise when one goes beyond a standard, thin, rotating disk model, for example by adding non-circular motions (Sect. 4.3.2) and/or extraplanar gas (Sect. 4.3.3). Degeneracies can be a concern even for high-resolution data with good signal-to-noise ratio (SNR), and becomes a very severe problem for barely resolved, low-SNR data. Ideally, next-generation modeling tools must be able not only to infer a particular galaxy property, but also to determine a realistic observational uncertainty. A robust assessment of the uncertainties, including potential covariances, is key to test DM models and modified-gravity theories.

In principle, the use of Bayesian inference to fully explore the parameter space should automatically account for all possible degeneracies, providing realistic uncertainties for all the parameters depending on their intrinsic correlation and on the quality of the data themselves. Therefore, it is natural to expect that Bayesian techniques will be adopted in the modeling of SKA-Mid HI data (but see below for a possible alternative). However, these techniques have two important drawbacks that we must overcome before considering them as the “golden standard”.

The first drawback is that, by construction, Bayesian inference is based on the likelihood function, which fully shapes the inferred posterior distributions, thus our measurements and associated uncertainties. The definition of this function is somewhat arbitrary and complicated by intrinsic difference between the model, which is smooth and typically axisymmetric, and the data, which show a higher degree of complexity especially at high resolutions, as well as by the presence of correlated noise in the datacube (e.g., [Davis et al., 2017a](#)).

The second drawback is the large computational cost needed for Bayesian analysis, which generally uses Monte Carlo methods and requires building $\sim 10^5 - 10^6$ models to robustly sample the parameter space. One way to optimize the existing algorithms is mapping the underlying probability distribution function and strategically sample it to improve both efficiency and accuracy. Possible avenues to do so include importance sampling, adaptive sampling, variance reduction techniques, optimal weighting, and multiphase Monte Carlo. Another, complementary way to optimize the process is to resort to GPU-acceleration, both for the creation of individual models and for the exploration of the parameter space. Libraries that simplify GPU-accelerated programming are now available in several programming languages (e.g., the JAX library in python, [Bradbury et al., 2018](#)), and can substantially boost the performances. Preliminary tests using KinMS ([Davis et al., 2013, 2020](#)) indicate a factor ~ 20 speed-up in the model creation.

4.4.2 Machine learning: artificial and Bayesian neural networks

An alternative approach to Bayesian fitting is to directly infer the desired quantities (e.g., flat rotation velocity, or DM density profile) directly from the 2D or 3D data using machine learning techniques, such as artificial neural networks (ANN). Neural networks have been shown to be able to detect and characterize sources from 2D or 3D HI data, notably as part of the first two SKA data challenges (Bonaldi et al., 2021; Hartley et al., 2023; Cornu et al., 2024, 2025). This approach would require generating mock data sampling different galaxy and DM parameters to create a large and representative training set. Once the ANN is trained, it should in principle provide fast predictions and hence be applicable to large datasets. However, the high dimensionality of the parameter space may induce degeneracies such that not all information may be recovered. In addition, ANNs generally lack explicit uncertainty quantification. Incorporating principles of Bayesian inference in a neural network, Bayesian neural networks (BNN) may constitute a promising way to device the new generation of dynamical modeling tools. BNNs indeed provide full posterior distributions over the model parameters, hence offering a way to quantify uncertainty. Furthermore, they may provide a more physically-tractable way to emulate current forward-modeling tools. Self-supervised, physics-aware, BNNs have been developed to fit gas kinematics in 2D (Dawson et al., 2021) and show significant promise when extended to 3D.

It may also be possible to adopt hybrid approaches, such as using an ANN for a fast initial parameter estimation, then a more computationally intensive Bayesian method to refine the results and quantify uncertainties. Both classical and Bayesian machine learning methods have been developed to yield dynamical masses (e.g. Ntampaka et al., 2015; Calderon and Berlind, 2019; Villanueva-Domingo et al., 2022; Chu et al., 2024; Hahn et al., 2024), although not specifically from HI data so far. Finally, since evaluating a rotation curve or a DM density profile from 2D or 3D data involves deconvolution steps, attention-based neural networks are likely to be well-suited to the task. Indeed, such neural network dynamically weigh the relevance of the different input elements, directly correlate all parts of the input, and provide increased interpretability through the attention weights. Furthermore, architectures such as transformers enable parallelization and hence faster training on modern hardware.

4.4.3 Kinematic modelling in the uv plane

Interferometric observations consist in a sparse sampling of the Fourier angular space (or uv -plane), providing the so-called visibilities. Conversion to the real space (the sky plane) requires a manipulation of the visibilities (e.g., gridding, weighting, deconvolution, and cleaning) that is not unique, and that inevitably alters the intrinsic properties of the original dataset. The ideal approach to circumvent these issues is to model the data directly in the uv -plane. This allows using the data in their original space, which provides two main advantages for modeling purposes. First, the uncertainties in the visibilities are well defined. Second, convolution with the synthesized beam (which is often the computational bottleneck in the model generation) is bypassed: the model and the data will probe exactly the same angular scales, by construction. This makes this method especially promising when dealing with barely resolved observations.

Several kinematic studies have explored this method (e.g., Dye et al., 2018; Rizzo et al., 2020;

Powell et al., 2021; Amvrosiadis et al., 2025), but the publicly available modeling packages have yet to implement it. Although promising, this method needs to be robustly implemented and carefully tested in order to assess its applicability to the unprecedented data volume that SKA-Mid will provide. Importantly, while the mathematical approach of a direct modeling of the visibilities is the most accurate one, it has two practical drawbacks: (1) the computational costs, and (2) the storage capacity costs. Regarding the first drawback, the computational cost can be large because, in a typical application, the disk model is first conceived in the real-space and only later transported to the uv plane with a Fast Fourier Transform (FFT) to be compared with the observed visibilities (e.g., Amvrosiadis et al., 2025). In this case, the original model must be created with sufficient resolution to account for the full range of possible scales probed by the data, which might be unpractical. This problem has been mitigated by applying a non-uniform FFT (Rizzo et al., 2020; Powell et al., 2021). The second drawback is a logistic one: the SKA observatory should store and provide the (raw or gridded) visibilities rather than more practical sub-cubes with the individual HI detections. An approach that would not require a storage of the visibilities, but only the dirty data cubes and the dirty beams, would be to work on the gridded visibilities instead of the measured ones. The fit quality would then be calculated in the Fourier domain by comparing a FFT of the dirty cube (done once in the complete process) with the FFT of the model multiplied with the gridded sampling function, i.e., the FFT of the dirty beam. This method has not been tested yet, but is theoretically even less computationally expensive than the 3D fitting method in the image domain, which requires two FFTs (back and forth) for the convolution with the clean beam. However, it would require the storage of the dirty cube and the dirty beam.

5 Conclusions

To be written

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